LYRA

the **L**arge-**Y**ield **Ra**diometer onboard PROBA2

Components of soft X-ray and extreme ultraviolet in flares observed by LYRA

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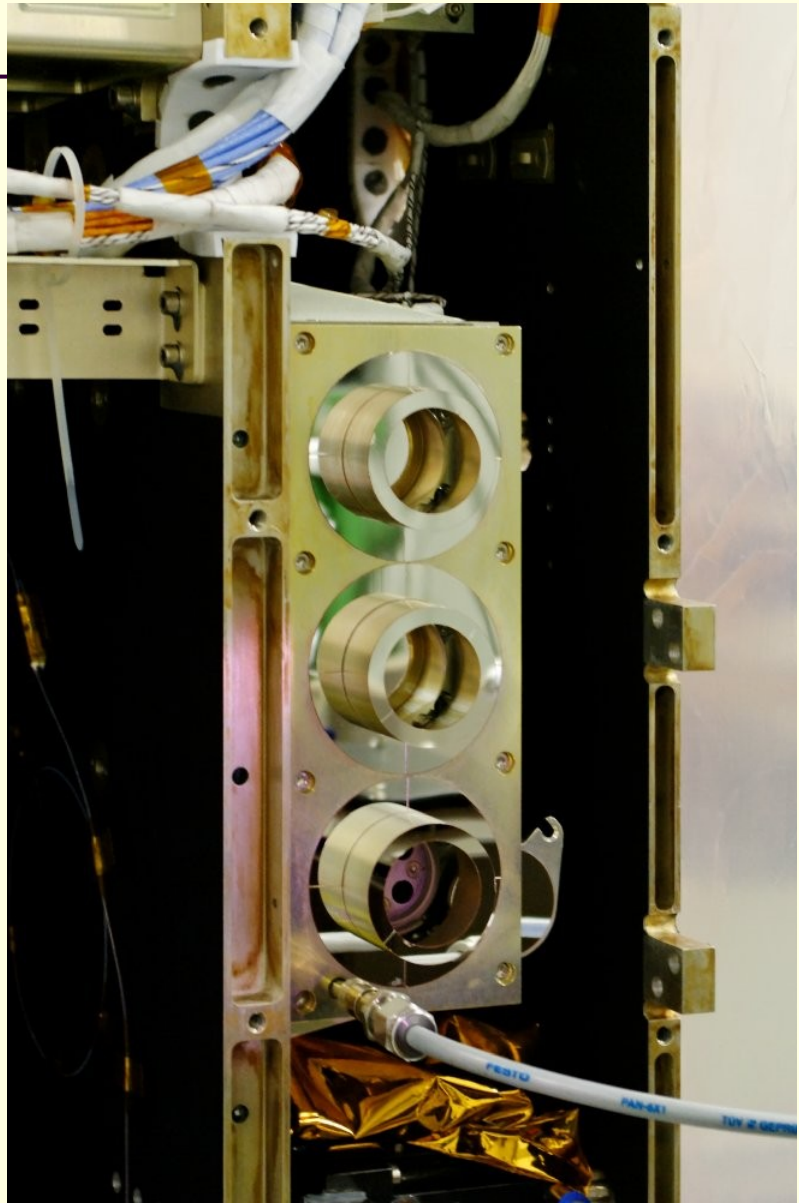
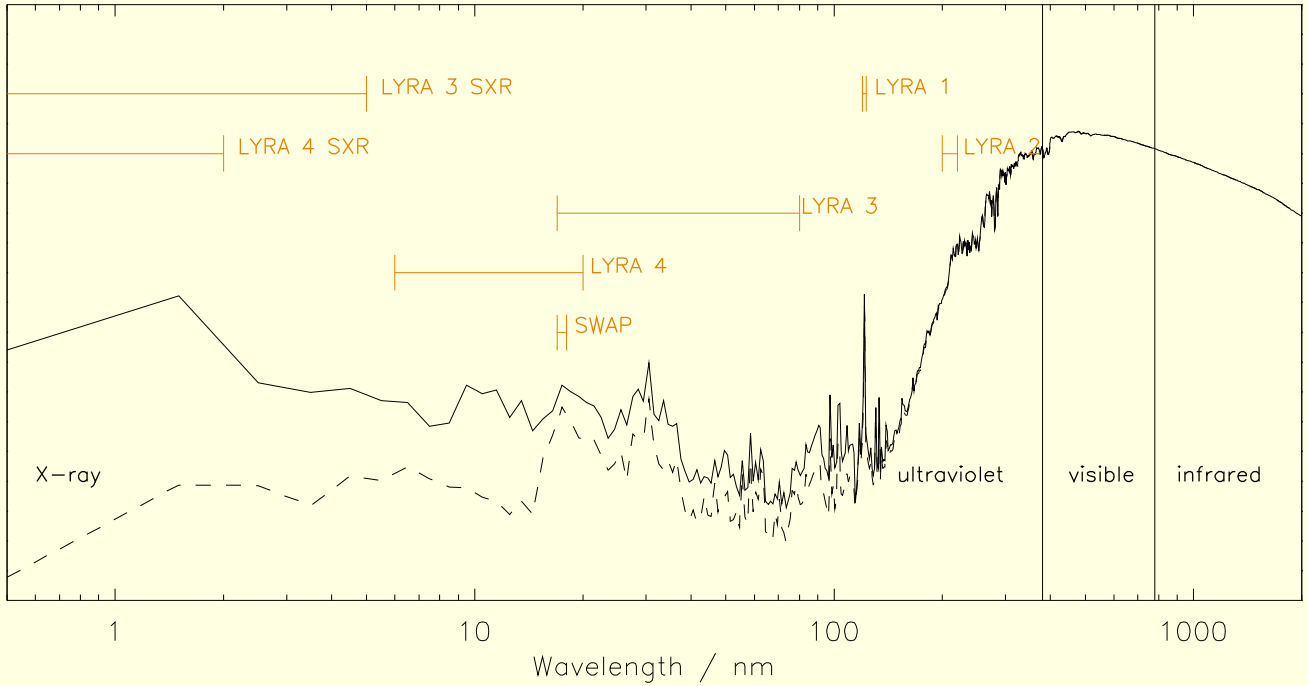
SCSL / PROBA2 Splinter Meeting
ESWW8, Namur, 29 Dec 2011

LYRA: the Large-Yield RAdiometer

- 3 instrument units (redundancy)
- 4 spectral channels per head
- 3 types of detectors,
  Silicon + 2 types of
  diamond detectors (MSM, PIN):
  - radiation resistant
  - insensitive to visible light
    compared to Si detectors
- High cadence up to 100 Hz

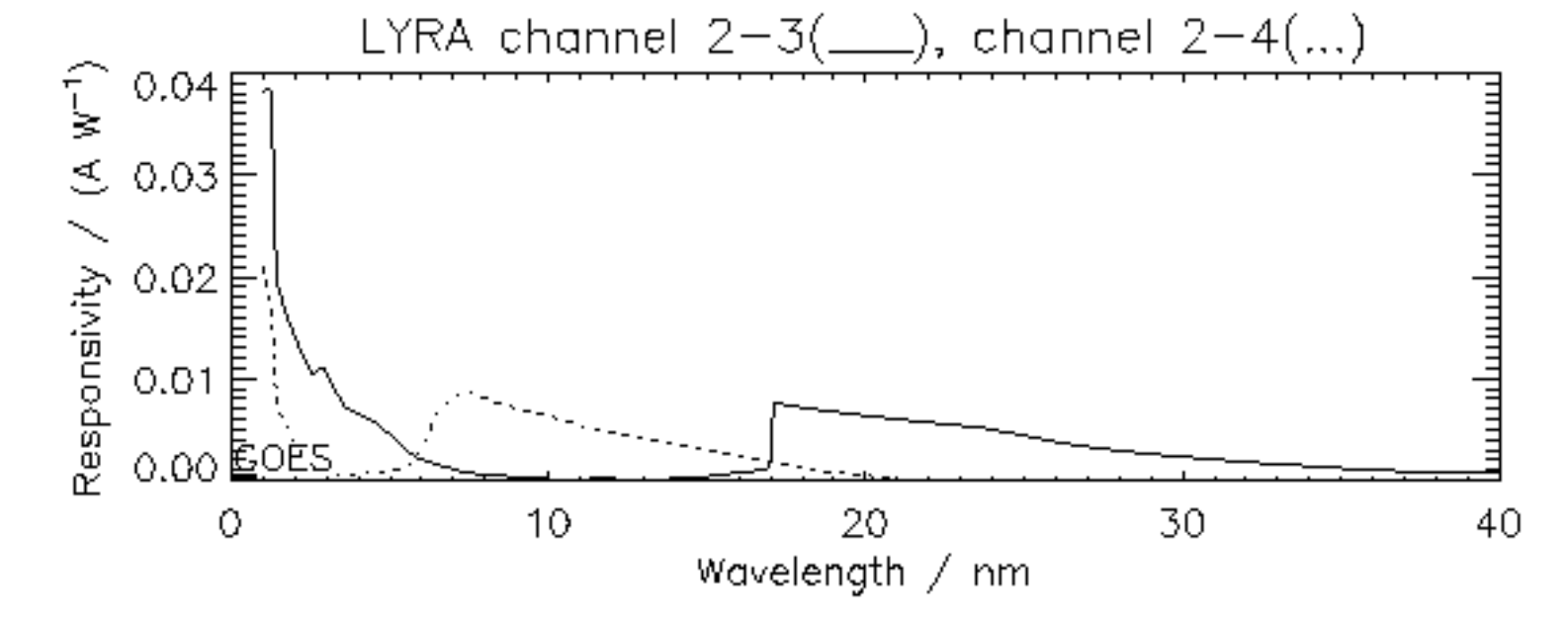


SWAP and LYRA spectral intervals for solar flares, space weather, and aeronomy



LYRA channel 1: the H I 121.6 nm Lyman-alpha line (120-123 nm)
LYRA channel 2: the 200-220 nm Herzberg continuum range (now 190-222 nm)
LYRA channel 3: the 17-80 nm Aluminium filter range incl the He II 30.4 nm line (+ <5nm X-ray)
LYRA channel 4: the 6-20 nm Zirconium filter range with highest solar variablility (+ <2nm X-ray)
SWAP: the range around 17.4 nm including coronal lines like Fe IX and Fe X

LYRA spectral response



- channel 2-3 (Aluminium filter)
- channel 2-4 (Zirconium filter)
- pre-launch calibration at BESSY
- additional SXR components <5 nm, <2 nm
- for comparison: GOES 0.1-0.8 nm

Data product definition

- (“Level 0”, telemetry from PROBA2, internal)
- Level 1 = full raw data (counts/ms)
- Level 2 = calibrated physical data (W/m²)
  (Caution: preliminary status. Require versioning.)
- Level 3 = processed products (e.g. averages)
- Level 4 = plots of products
- Level 5 = event lists (optionally with plots)

New (well, more or less new) LYRA products

… resulting from calibration attempts:

- Level 2 FITS files
- Level 3 FITS files
- (Level 4) One-day overviews
- (Level 4) Three-day overviews
- (Level 5) Flare lists
- (Level 5) GOES vs. LYRA proxies (preliminary)

… available here at the P2SC website:

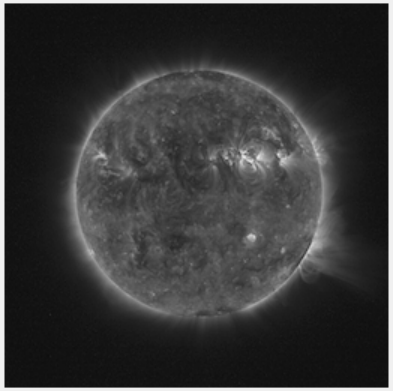
http://proba2.sidc.be/

PROBA2 SCIENCE CENTER

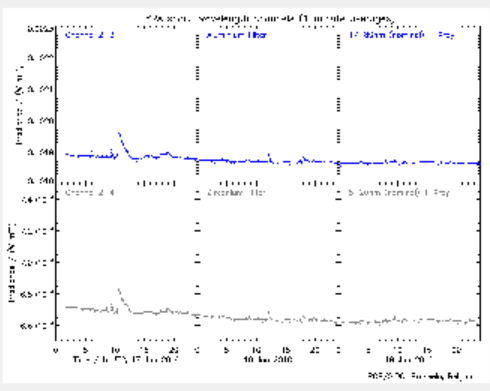
Last update: 18th of July 2011

Welcome to the PROBA2 science center.

Check out the news on the July 1st solar eclipse and the June 7 M-flare!



Watch the latest SWAP image or movie



Go to the latest 3-day LYRA curve and quicklook daily image

Direct link to the scientific data:

- SWAP calibrated FITS - SWAP movies
- LYRA calibrated FITS - LYRA quicklook PNGs over a day and over 3 days - LYRA flare list

More info on all available data here. Consult the tutorial on SWAP analysis software.

Data gaps or unusual data? Please consult the Google Calendars for SWAP and LYRA to find out about special campaigns and off-pointing sequences (commanded via SWAP but also affecting LYRA signals).

NEW : LYRA calibrated data available on-line + an interactive Quicklook Viewer.

Contact us via swap_lyra *at* sidc.be.

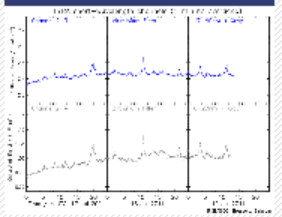
Latest news

2 July 2011
Partial solar eclipse of July 1, 2011

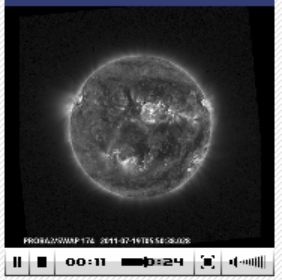
7 June 2011
Extra-ordinary M-flare observations

2 June 2011
June 1st partial solar eclipse

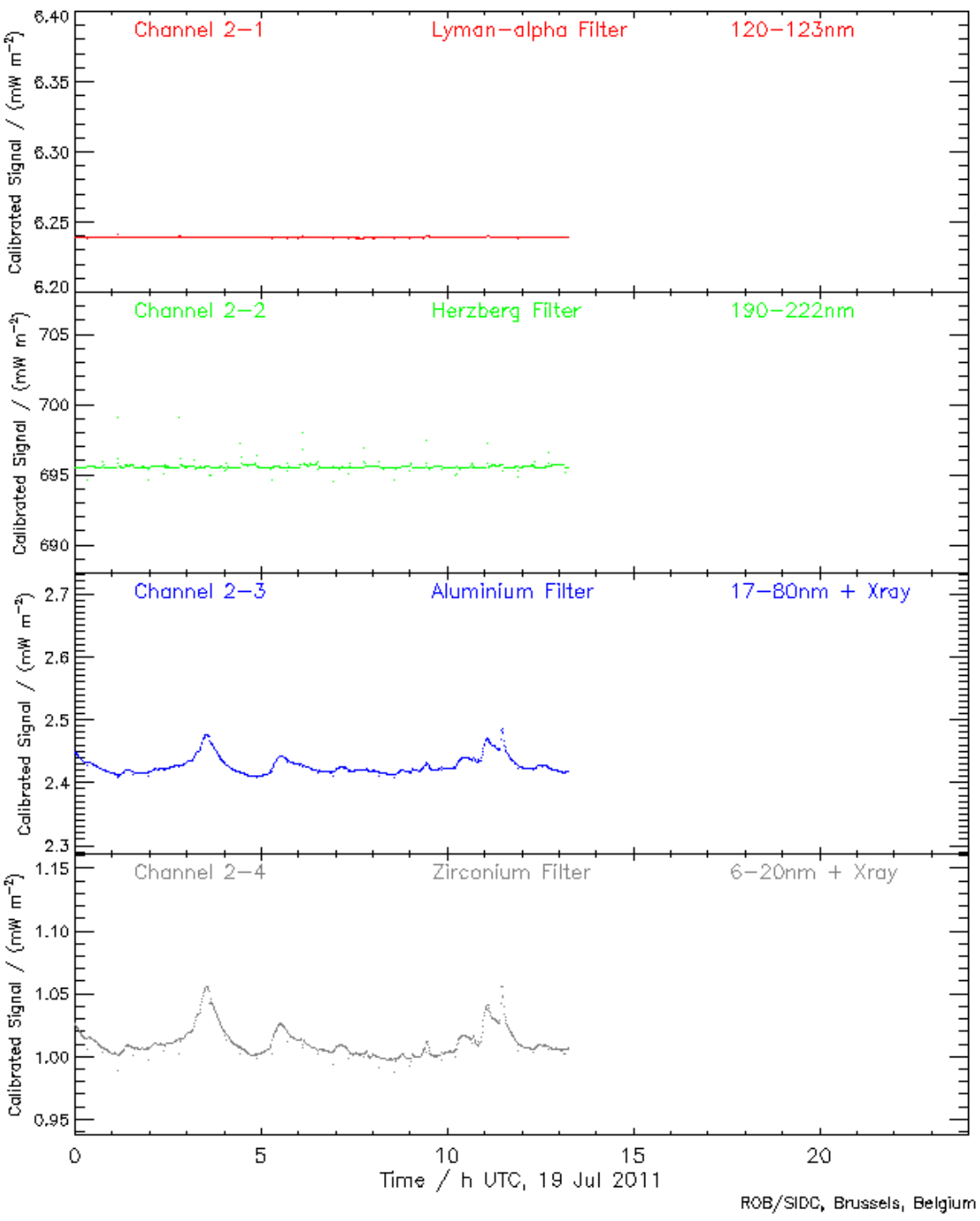
Latest LYRA curve



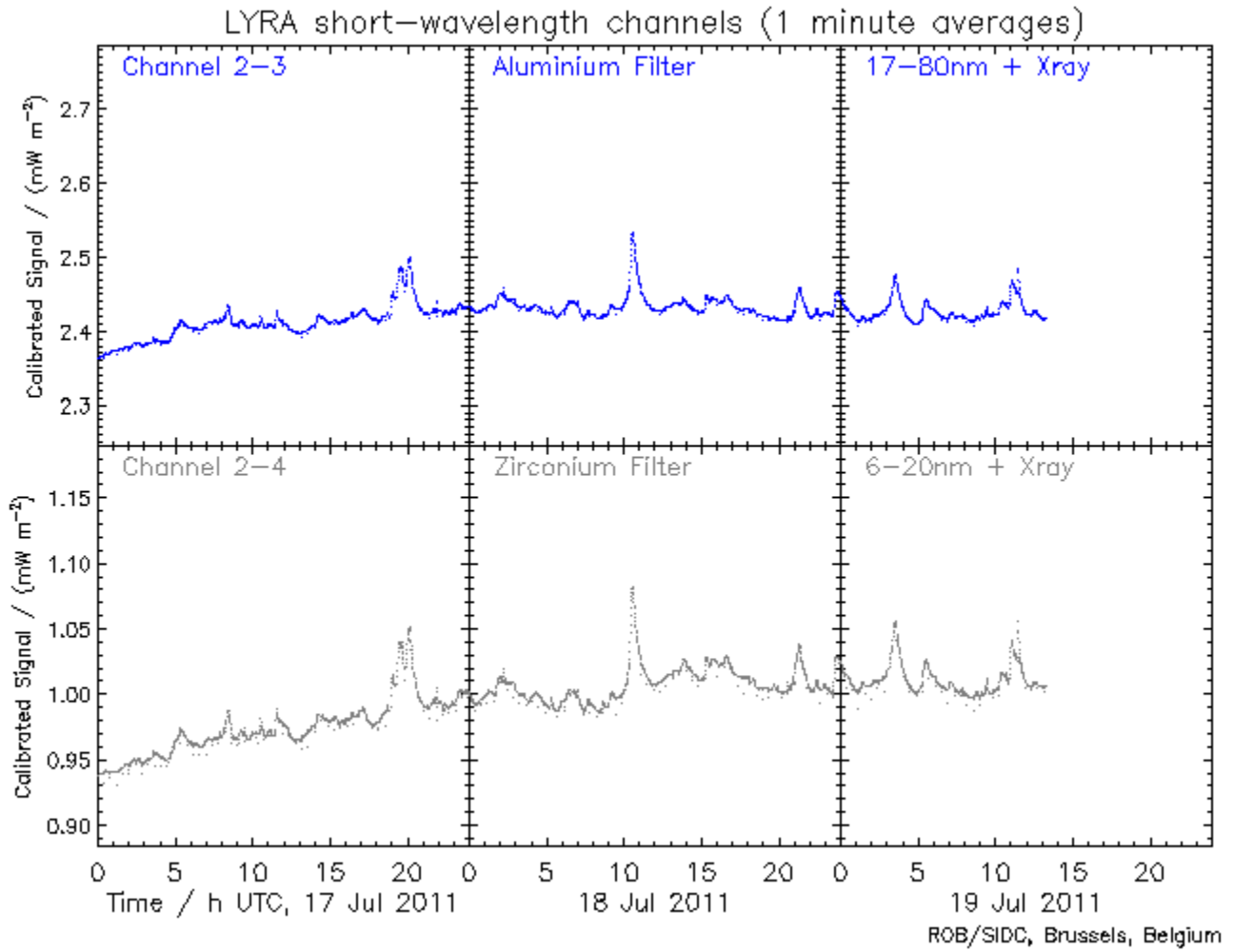
Latest SWAP movie



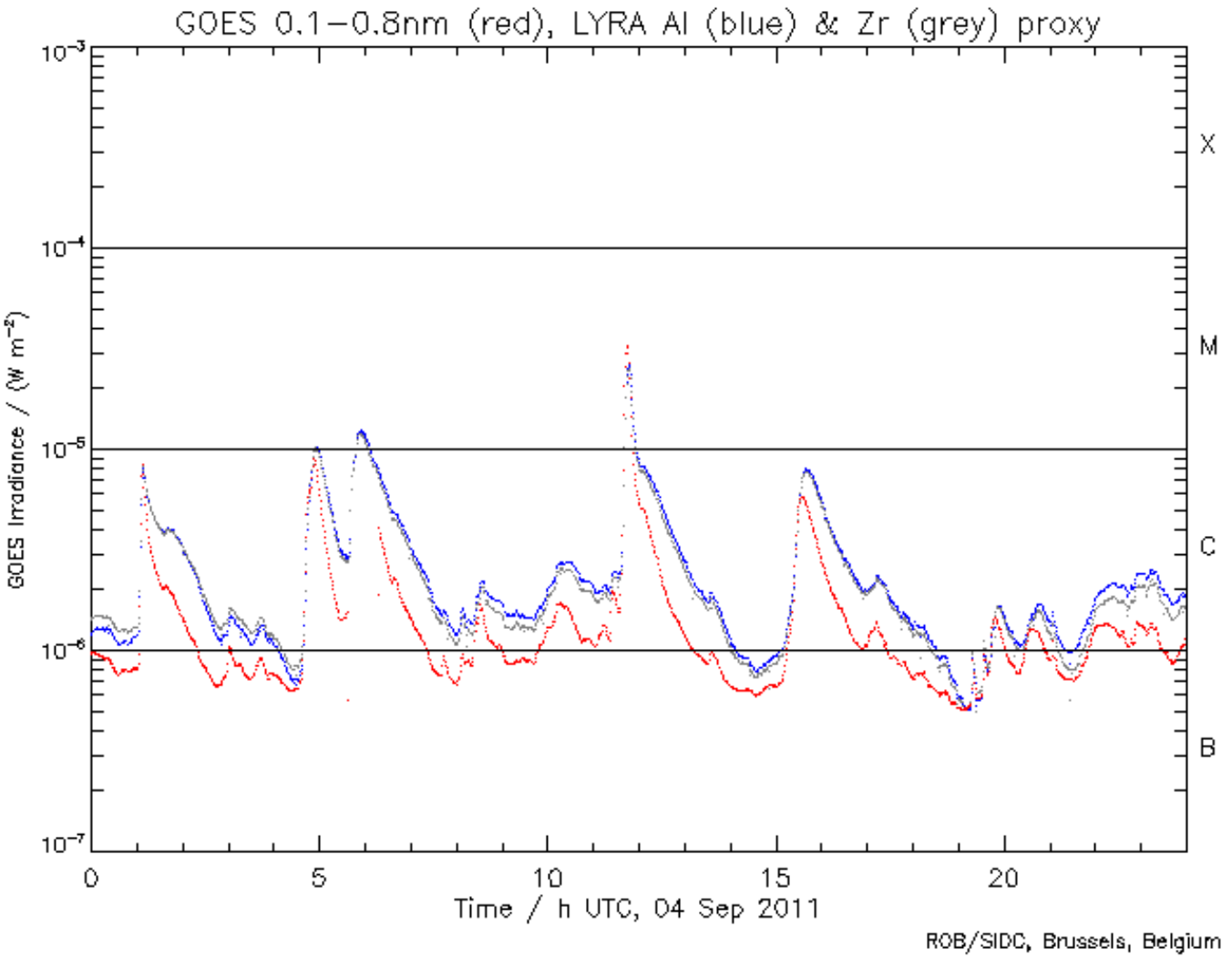
one-day overview



three-day overview



GOES vs. LYRA proxies



For more information on this subject, please see the talk by David Berghmans,
Automated flare detection and localization with PROBA2
Session 4A, Thursday 01 December 2011, 12:20

2011 LYRA Flare List

Please note:

- This list uses the "G14" and "G15" X-ray entries from the "Edited Events" lists of the NOAA Space Weather Prediction Center.
- The purpose of the list is to get an overview of the flares that LYRA observes and relate them to class, begin, max, etc. according to GOES.
- In the daily images (follow links below), the flares are marked at the top of the LYRA Zr-channel curve, with event number and class corresponding to their temporal "begin".
- Images of intervals around each flare (1h before, 2h after) are linked to a list below this curve (follow event links), given that LYRA was observing during this interval.
- In the flare images, event number and class are again marked at the top, corresponding to their "begin", while the GOES maximum is marked with a short vertical line. These images contain all four LYRA channels, plus one GOES channel.
- From November to January, PROBA2 experiences "eclipse season". For several minutes during each orbit, the solar disk is occulted by the Earth, and the observed irradiances decrease to dark-current levels - thus LYRA flares may be (partially) hidden.
- Additional information can be found at the PROBA2 website on the LYRA daily quicklook page or on the LYRA 3day quicklook page.

2010 page

July 2011

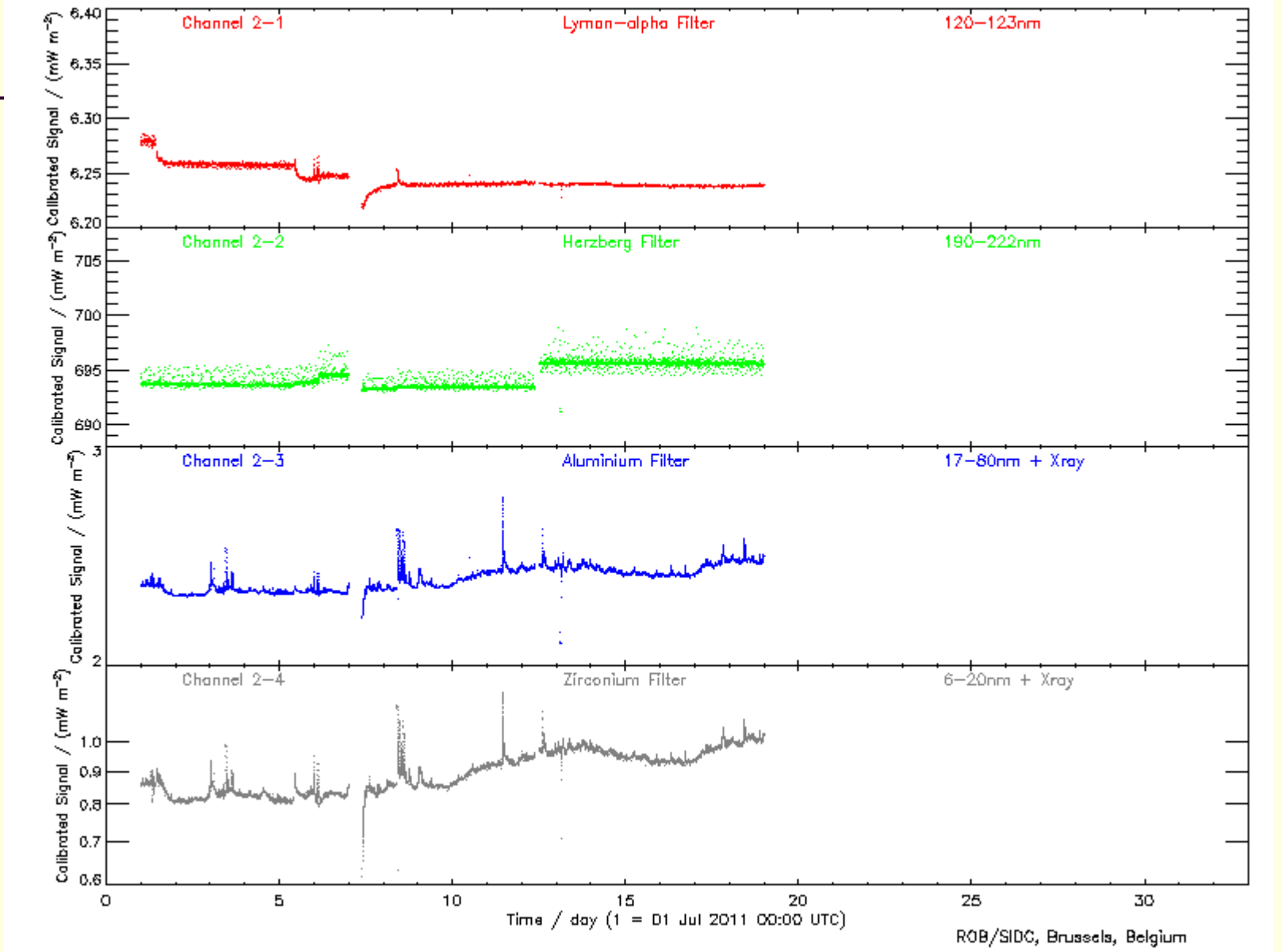
				Fri 01	Sat 02	Sun 03
Mon 04	Tue 05	Wed 06	Thu 07	Fri 08	Sat 09	Sun 10
Mon 11	Tue 12	Wed 13	Thu 14	Fri 15	Sat 16	Sun 17
Mon 18	Tue 19	Wed 20	Thu 21	Fri 22	Sat 23	Sun 24
Mon 25	Tue 26	Wed 27	Thu 28	Fri 29	Sat 30	Sun 31

June 2011

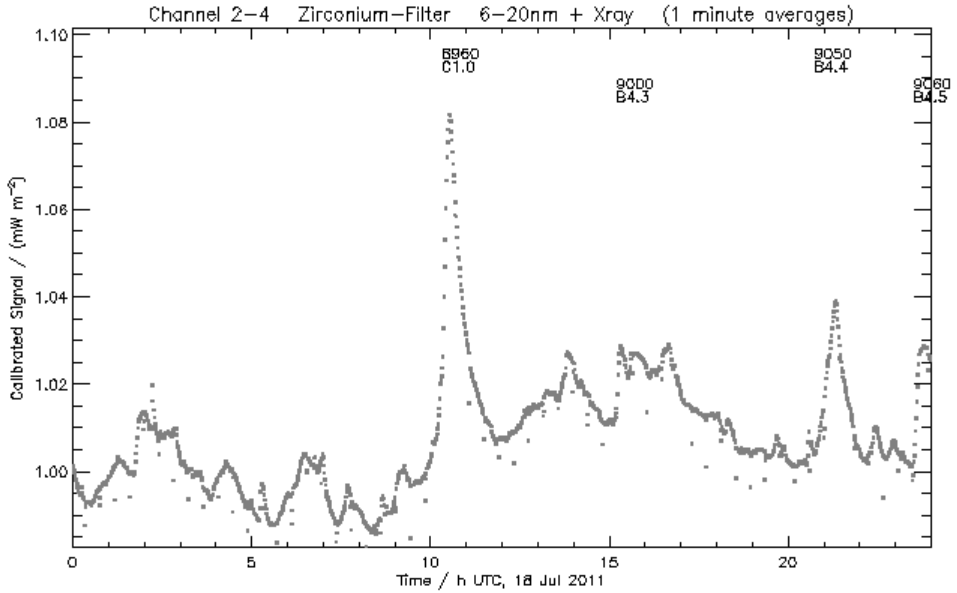
		Wed 01	Thu 02	Fri 03	Sat 04	Sun 05
Mon 06	Tue 07	Wed 08	Thu 09	Fri 10	Sat 11	Sun 12
Mon 13	Tue 14	Wed 15	Thu 16	Fri 17	Sat 18	Sun 19
Mon 20	Tue 21	Wed 22	Thu 23	Fri 24	Sat 25	Sun 26
Mon 27	Tue 28	Wed 29	Thu 30			

May 2011

monthly overview



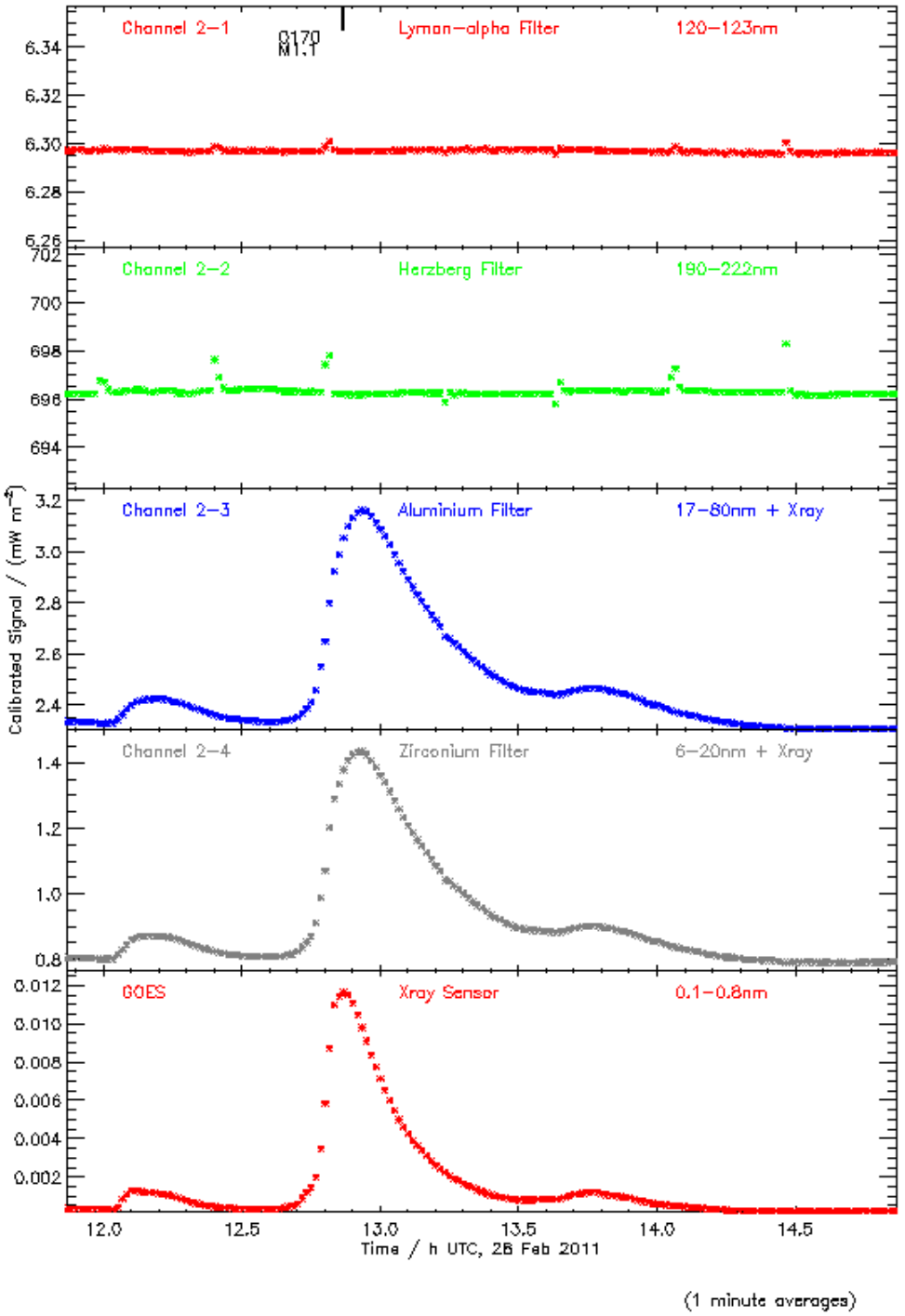
18 Jul 2011 Flare List



event	begin	max	end	class	region
8960	10:19	10:28	10:38	C1.0	1254
9000	15:12	15:17	15:24	B4.3	1254
9050	20:44	21:14	21:29	B4.4	
9060	23:31	23:40	23:58	B4.5	1255

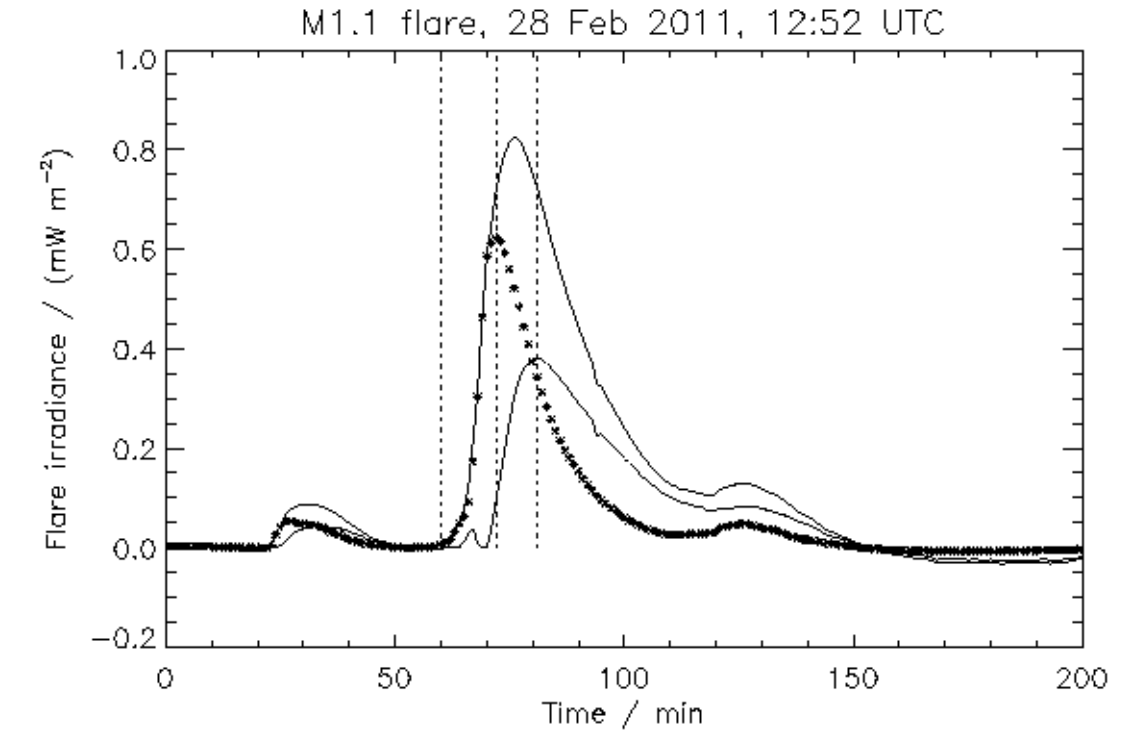
Example:
M1.1 flare, 28 Feb 2011

- start to rise at same time
- parallel in impulsive phase
- GOES peaks earlier
- LYRA decreases slower
- linear factor in pure flare irradiance



Flare components

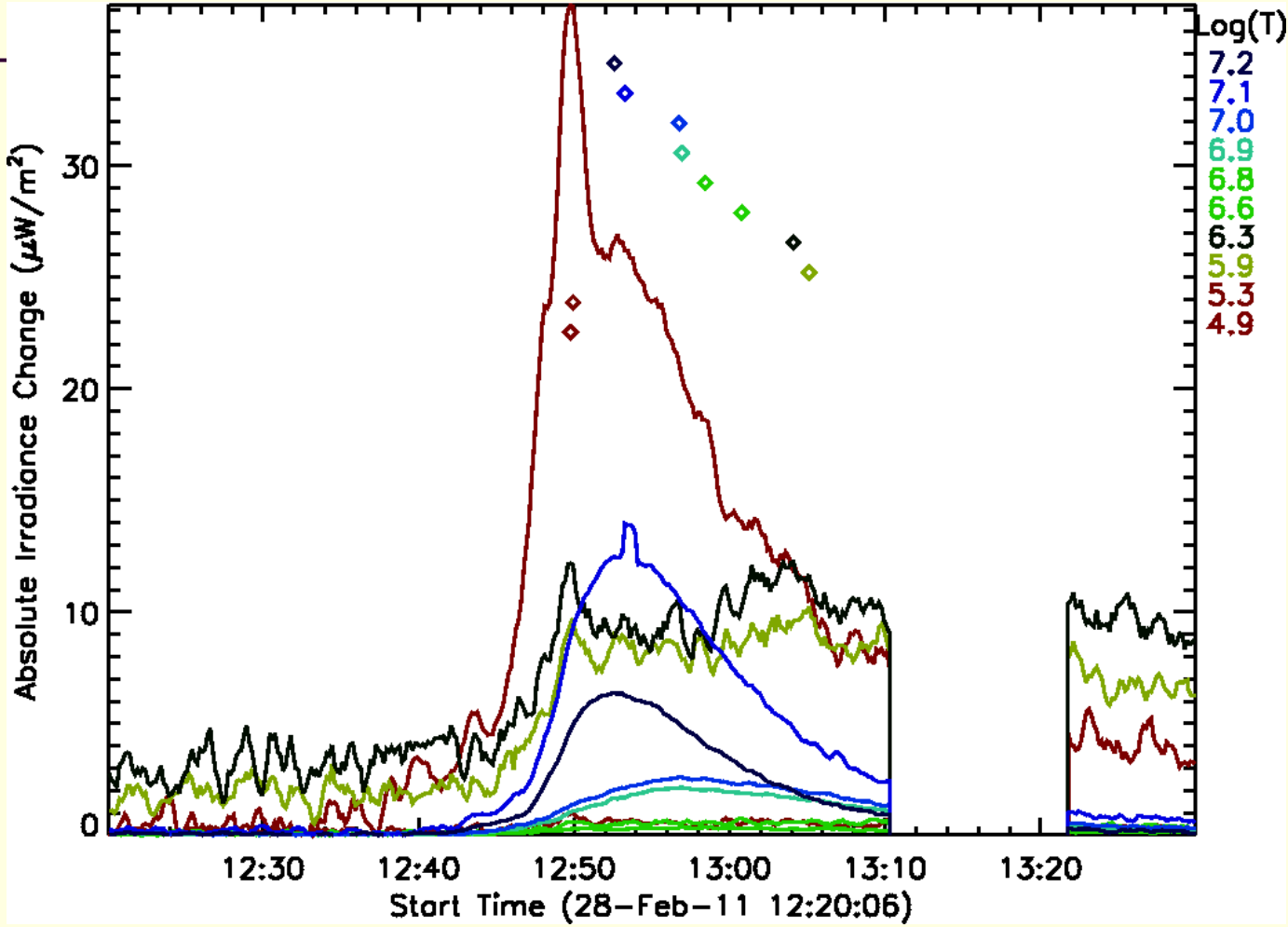
ch2-3 = SXR+EUV



- “SXR”: emission with log(T)>7
- “EUV residual”: emission with 6<log(T)<7
- “little bump”: emission with log(T)<6

Compare with SDO/EVE:

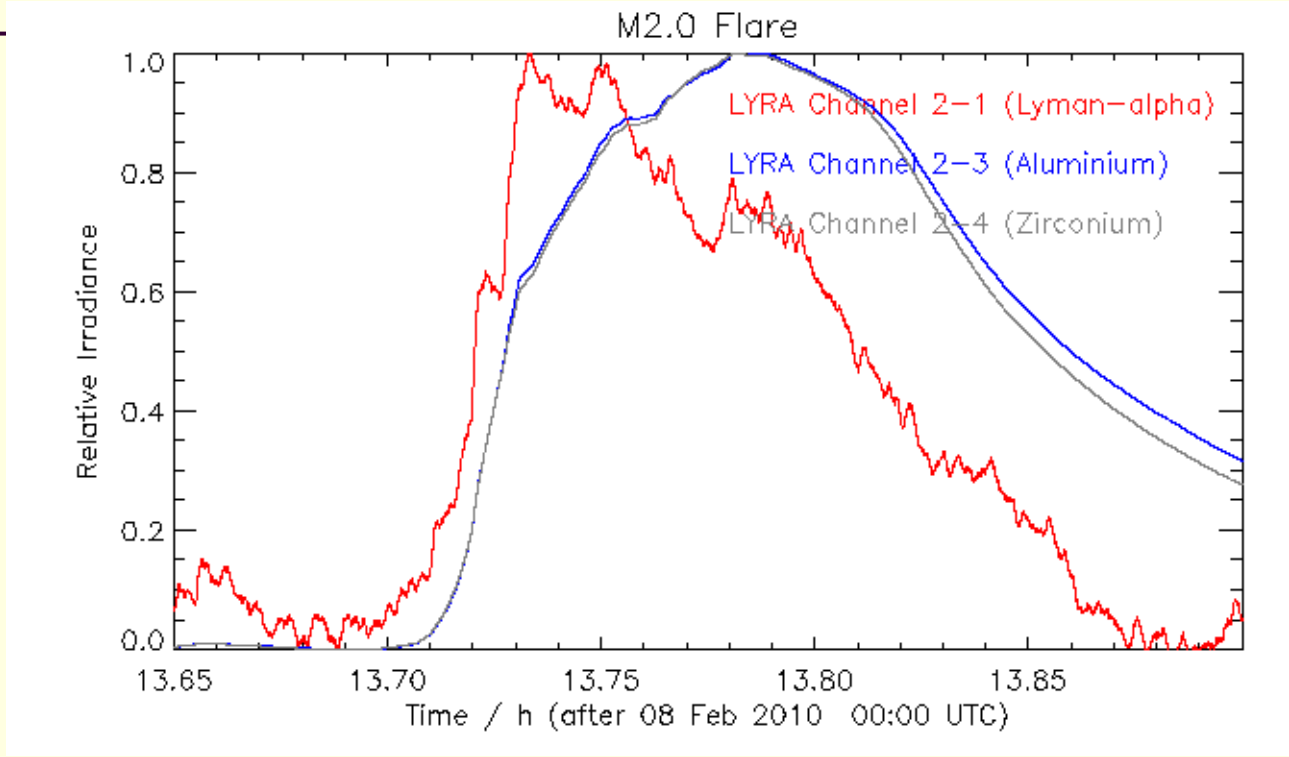
Thermal evolution plot



based on:

- solar spectra observed by SDO/EVE
- contribution functions from the CHIANTI atomic database

Lyman-alpha signal



- LYRA in early 2010
- signal peaks in rising phase
- log(T)<6

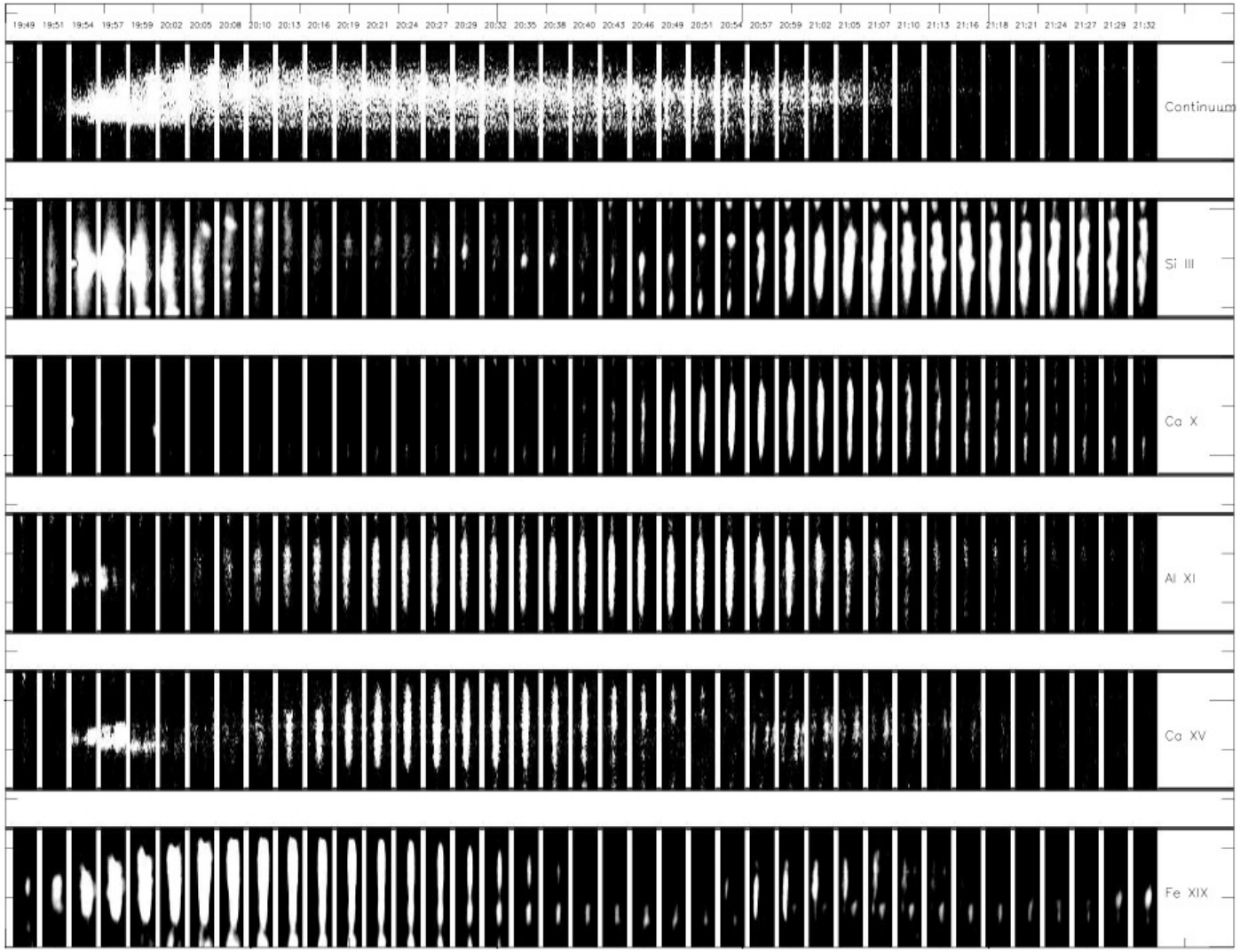


FIG. 8.—Background-subtracted spectra for the free-free continuum, for Si III, Ca X, Al XI, Ca XV, and Fe XIX, as a function of time. A Ne VI second-order line is present in the Fe XIX spectral window between 20:50 and 21:10 UT.

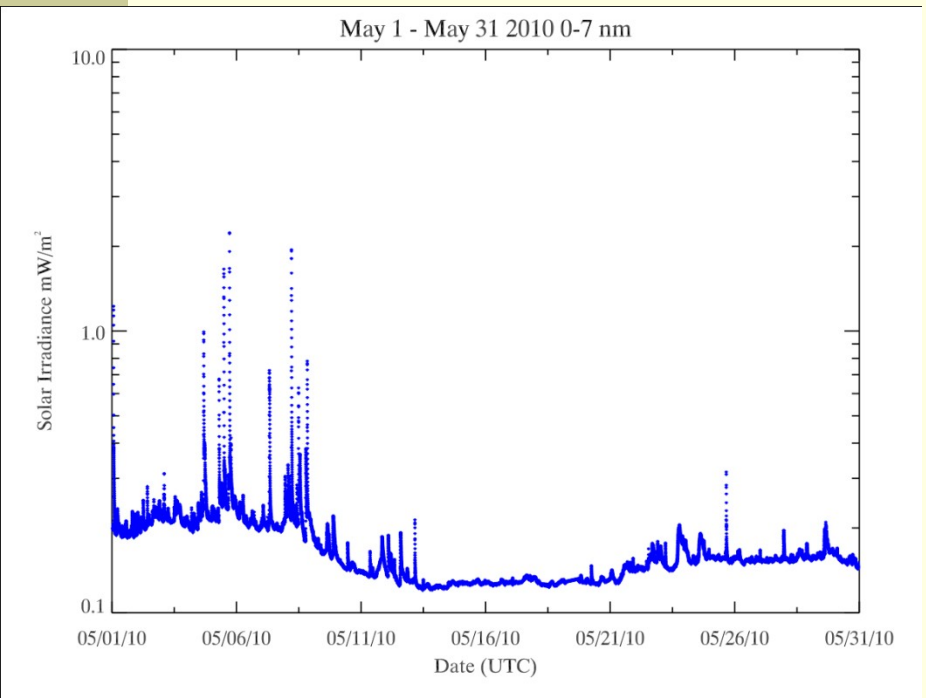
Why does it matter?

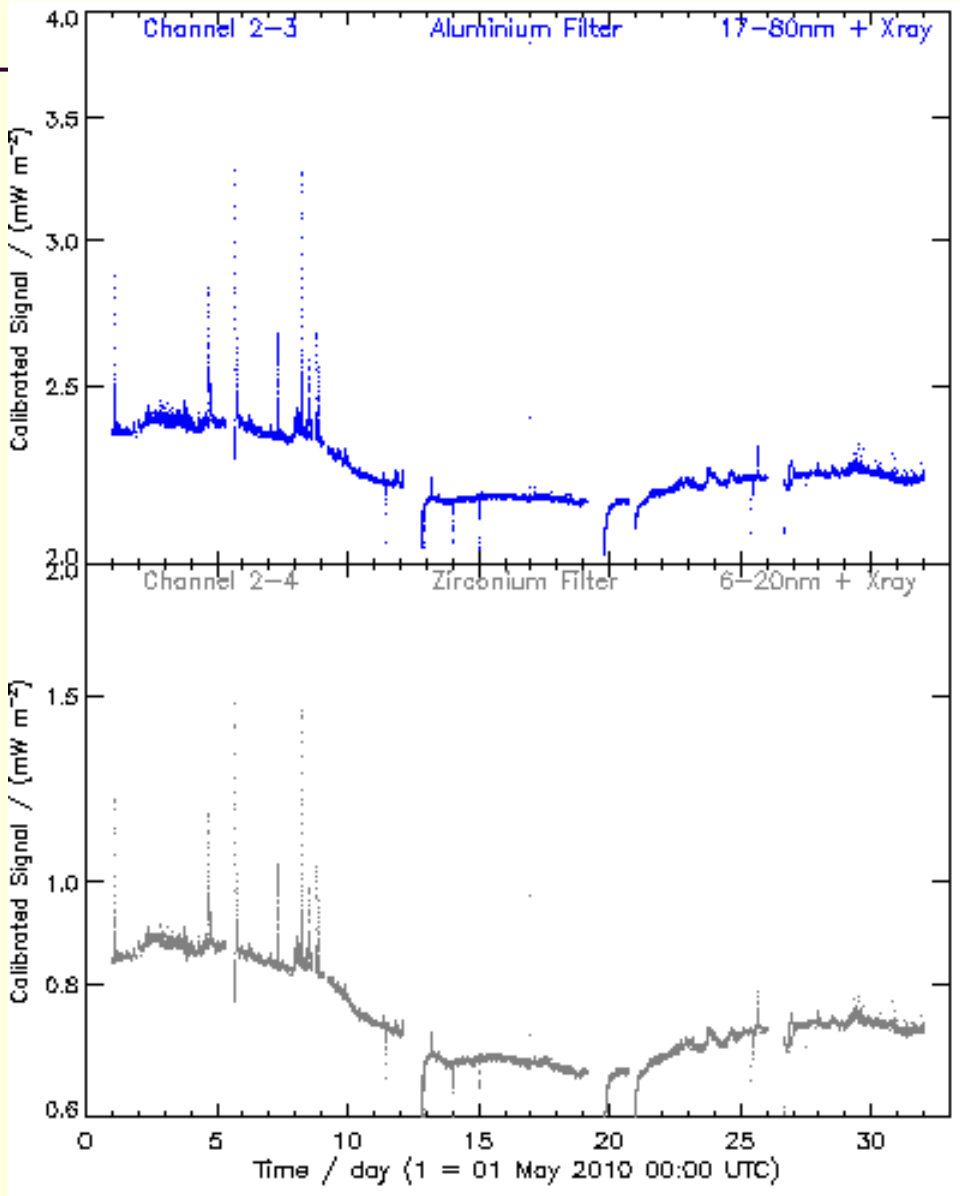
- SXR component, ~0.1nm - ?nm, mostly optically thin emission lines, effect on E region ionosphere, weakly dependent on flare location
- EUV component, ?nm - ~120nm, many important lines optically thick, effect on F region ionosphere, dependent on flare location
- Model calculations and X flare observations (e.g. 2003) show: Solar flare enhancements depend on disk location: Disk-center flares have stronger effect on ionosphere than limb flares due to EUV enhancements
- Thermospheric response ~20 minutes faster for disk-center flare

- Satellite drag, radio communications, GPS accuracy: related to thermosphere and ionosphere conditions => location and SXR/EUV components determine space weather consequences of flare effects

  L. Qian et al., "Flare location on the solar disk: Modeling the thermosphere and ionosphere response", Journal of Geophysical Research 115, A09311 (2010)

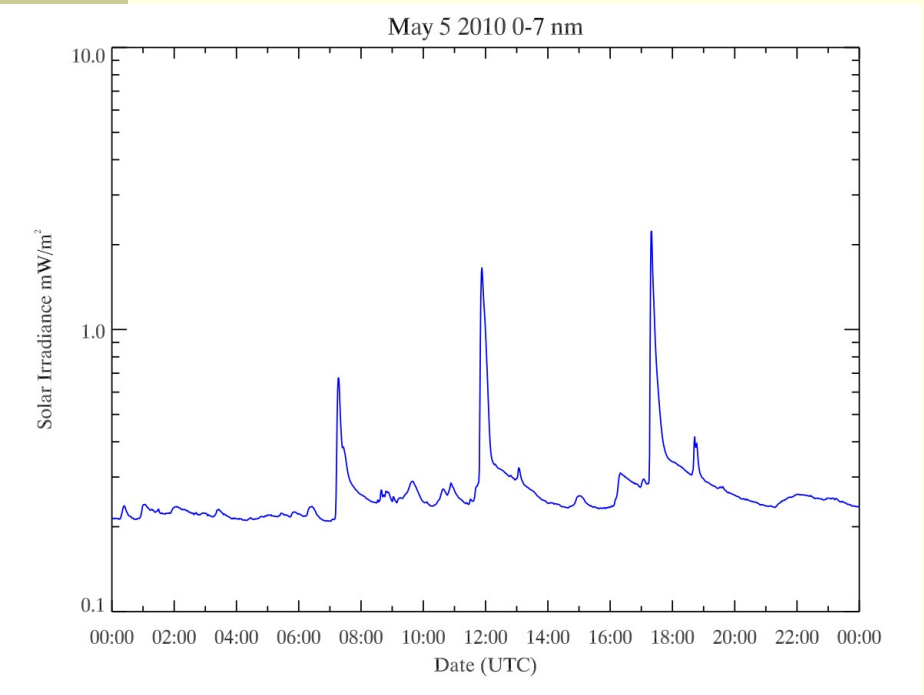
May 2010 EVE vs. LYRA

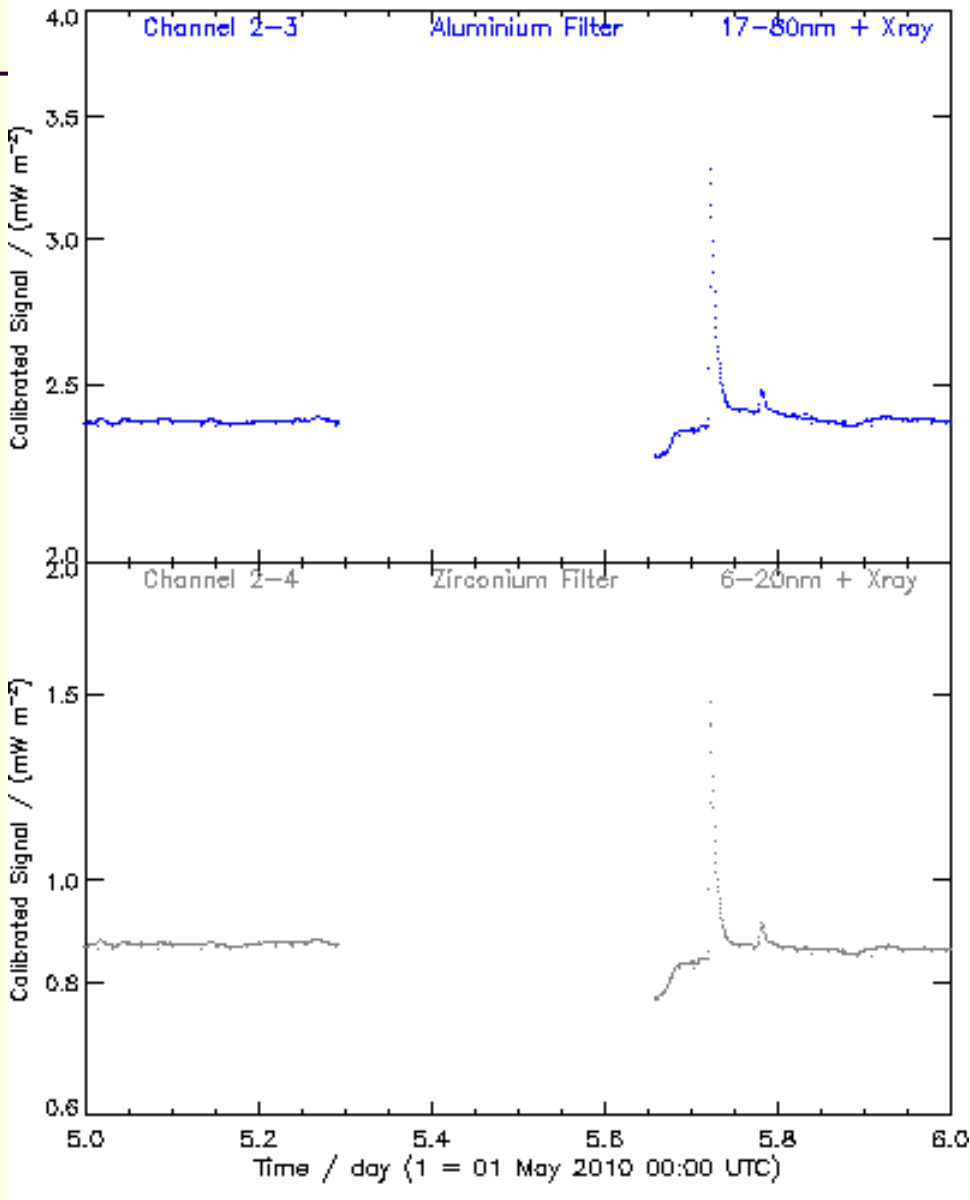




M1.2 flare 05 May 2010 17:19 UTC

GOES (0.1 - 0.8 nm) ~0.012 mW/m^2
LYRA (0 - 2 nm) ~0.7 mW/m^2
LYRA (0 - 5 nm) ~1.0 mW/m^2
EVE (0 - 7 nm) ~2.0 mW/m^2





Feb/Mar 2011 EVE vs. LYRA